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The SAURON project – XIII. SAURON–GALEX study of early-type galaxies: the ultraviolet colour–magnitude relations and Fundamental Planes

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ABSTRACT

We present *Galaxy Evolution Explorer* far-ultraviolet (FUV) and near-ultraviolet (NUV) imaging of 34 nearby early-type galaxies from the SAURON representative sample of 48 E/S0 galaxies, all of which have ground-based optical imaging from the MDM Observatory. The surface brightness profiles of nine galaxies (≈ 26 per cent) show regions with blue UV–optical colours suggesting RSF. Five of these (≈ 15 per cent) show blue integrated UV–optical colours that set them aside in the NUV integrated colour–magnitude relation. These are objects with either exceptionally intense and localized NUV fluxes or blue UV–optical colours throughout. They also have other properties confirming they have had RSF, in particular $H\beta$ absorption higher than expected for a quiescent population and a higher CO detection rate. This suggests that residual star formation is more common in early-type galaxies than we are used to believe. NUV blue galaxies are generally drawn from the lower stellar velocity dispersion ($\sigma_e < 200 \text{ km s}^{-1}$) and thus lower dynamical mass part of the sample. We have also constructed the first UV Fundamental Planes and show that NUV blue galaxies bias the slopes and increase the scatters. If they are eliminated, the fits get closer to expectations from the virial theorem. Although our analysis is based on a limited sample, it seems that a dominant fraction of the tilt and scatter of the UV Fundamental Planes is due to the presence of young stars in preferentially low-mass early-type galaxies. Interestingly, the UV–optical radial colour profiles reveal a variety of behaviours, with many galaxies showing signs of RSF, a central UV-upturn phenomenon, smooth but large-scale age and metallicity gradients and in many cases a

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combination of these. In addition, FUV–NUV and FUV–V colours even bluer than those normally associated with UV-upturn galaxies are observed at the centre of some quiescent galaxies. Four out of the five UV-upturn galaxies are slow rotators. These objects should thus pose interesting challenges to stellar evolutionary models of the UV upturn.

Key words: galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: fundamental parameters – galaxies: photometry – galaxies: structure – ultraviolet: galaxies.

1 INTRODUCTION

The dominant formation mechanism of early-type galaxies remains one of the long-standing debates of modern astrophysics. The classical monolithic model (e.g. Eggen, Lynden-Bell & Sandage 1962; Larson 1974) suggests that early-type galaxies form in highly efficient starbursts at high redshifts and evolve without much residual star formation from that point onward. On the other hand, the popular Λ cold dark matter (Λ CDM) paradigm strongly suggests a hierarchical merger scenario for massive early-type galaxies (e.g. Toomre & Toomre 1972). In this model, early-type galaxies form as the result of successive mergers and are thought to have continued or episodic star formation events.

Numerous tests discriminate between these scenarios. For example, the optical colour–magnitude relations (CMRs) of early-type galaxies consistently reveal a small scatter around the mean relation (e.g. Bower, Lucey & Ellis 1992; Ellis et al. 1997; Stanford, Eisenhardt & Dickinson 1998; van Dokkum et al. 2000), apparently supporting the monolithic scenario. Deep imaging surveys have, however, shown that many early-type galaxies possess shells, tidal features (e.g. Schweizer & Seitzer 1992) and signatures of ongoing or recent star formation (RSF; e.g. NGC 2865, Hau, Carter & Balcells 1999; Raimann et al. 2005). Integral-field spectroscopy now also allows us to obtain spatially resolved maps of various galaxy properties. For example, we have previously highlighted the importance of central stellar discs and kinematically decoupled cores in the SAURON sample (e.g. de Zeeuw et al. 2002, hereafter Paper II; Krajnović et al. 2008, hereafter Paper XII).

More recently, the *Galaxy Evolution Explorer* (GALEX) satellite opened up the ultraviolet (UV) window, allowing to probe the RSF history of galaxies with much greater accuracy than was previously possible with only optical information. The rest-frame UV is one order of magnitude more sensitive than the optical to the presence of hot stars, easily revealing populations younger than 1 Gyr. Yi et al. (2005) used GALEX Medium Imaging Survey data to construct the first near-ultraviolet (NUV) CMR of early-type galaxies classified by the Sloan Digital Sky Survey (SDSS). They found a remarkably large scatter towards blue colours and interpreted it as evidence of RSF. Kaviraj et al. (2007) and Schawinski et al. (2007) subsequently found that the fraction of early-type galaxies that experienced RSF in the last billion years can be greater than 30 per cent. Young, Bendo & Lucero (2009) further concluded that the star formation rates of nearby early-type galaxies derived from mid-infrared imaging are in good agreement with the UV results. There is thus clear evidence for residual star formation in the local early-type galaxy population and for its influence on global scaling relations, at least CMRs.

The Fundamental Plane (FP) is another key scaling relation of early-type galaxies, a two-dimensional plane in the three-dimensional manifold of their global structural parameters (effective radius, surface brightness and stellar velocity dispersion; e.g. Dressler et al. 1987; Djorgovski & Davis 1987), and may also

be affected by RSF. Under the assumption of structural homology (i.e. all early-type galaxies have the same mass distribution and kinematics), the virial theorem implies that the FP parameters should scale in a specific manner, but observations reveal a *tilt* of the FP away from the virial prediction. Numerous studies have investigated this, and the observed tilt has been attributed to a variation of the mass-to-light ratio across the sequence of early-type galaxies and/or to the breaking of the homology assumption (see e.g. Dressler et al. 1987; Djorgovski & Davis 1987; Guzmán, Lucey & Bower 1993; Jørgensen, Franx & Kjørgaard 1996; Scodreggio 1997; Pahre, Djorgovski & De Carvalho 1998a; Kelson et al. 2000; Bernardi et al. 2003; Treu et al. 2006; Cappellari et al. 2006, hereafter Paper IV; Bolton et al. 2008). Nevertheless, much debate still exists on the origin of the tilt and the scatter about the mean relation (e.g. D’Onofrio et al. 2006 and references therein).

Spatially resolved galaxy data provide valuable information on the details of the stellar population distributions, especially when combined with UV data. Hence, we aim here to obtain a UV data base matching the SAURON integral-field data, and present spatially resolved UV and optical imaging of most early-type galaxies in the SAURON survey. Numerous investigations can be performed with such a data base, but we will focus here on revisiting the effects of star formation on the scaling relations of early-type galaxies, particularly CMRs and the FP. Longstanding enigmas such as the UV-upturn phenomenon (see O’Connell 1999 and references therein) are closely related.

In this paper, we thus present and discuss our UV imaging of SAURON early-type galaxies obtained with GALEX, along with ground-based optical imaging from the MDM Observatory. In Section 2, we present a brief summary of the SAURON survey and our optical observations, and then describe at length the GALEX observations and surface photometry. In Section 3, we discuss the UV CMRs, the first UV FPs, and make an attempt at interpreting radial UV–optical colour profiles. We summarize our findings in Section 4.

2 OBSERVATIONS AND DATA REDUCTION

2.1 SAURON integral-field spectroscopy

The SAURON observations are aimed at determining the two-dimensional stellar kinematics (Emsellem et al. 2004, hereafter Paper III), stellar line strengths (Kuntschner et al. 2006, hereafter Paper VI) and ionized gas kinematics (Sarzi et al. 2006, hereafter Paper V) of 48 nearby early-type galaxies and 24 spiral bulges in the field and clusters (see Paper II). The observations and data reduction are described in the relevant papers. Detailed dynamical modelling is also available (see e.g. Paper IV; Cappellari et al. 2007, hereafter Paper X; Emsellem et al. 2007, hereafter Paper IX).

2.2 Optical observations

Ground-based optical imaging observations in the *Hubble Space Telescope* (HST) filter F555W (similar to Johnson *V*) were obtained using the MDM Observatory 1.3-m McGraw-Hill Telescope, as part of a large survey targeting the whole SAURON galaxy sample. The MDM observations are described in detail in Falc3n-Barroso et al. (in preparation) and were reduced and calibrated in the standard manner. The field of view of the MDM images is 17.3×17.3 arcmin² with 0.508×0.508 arcsec² pixels, allowing for accurate sky subtraction and proper sampling of the seeing. The seeing during the observations was typically ≈ 1.2 arcsec. The total *V* magnitudes we use in this paper are also from Falc3n-Barroso et al. (in preparation).

2.3 UV observations

We obtained both FUV (1350–1750 Å) and NUV (1750–2750 Å) images of 34 E/S0 galaxies using the medium-depth imaging mode of *GALEX*, as part of an ongoing UV imaging survey of the SAURON sample (*GALEX* guest investigator programmes GI1–109 and GI3–041) and the *GALEX* Nearby Galaxy Survey (Gil de Paz et al. 2007). Of the 48 early-type galaxies in the SAURON sample, six are too close to UV-bright stars to be observed with *GALEX* and three from our programmes have not been observed yet, similarly for five galaxies from other guest investigator programmes, explaining the current sample of 34. The typical exposure time per field was one orbit (≈ 1700 s) but occasionally more. Details of the *GALEX* instruments, pipeline and calibration are described in Martin et al. (2005) and Morrissey et al. (2005, 2007). We note in particular a possible systematic error in the FUV and NUV zero-points of up to 0.15 mag. This is not included in the uncertainties quoted in the current paper as we are mostly interested in the relative colours of our objects. The spatial resolutions of the images are approximately 4.5 and 6.0 arcsec full width at half-maximum in FUV and NUV, respectively, sampled with 1.5×1.5 arcsec² pixels.

2.4 Photometry

The *GALEX* UV images are delivered pre-processed but we undertook our own estimate of the sky values. This was calculated for each image as the mean of the sky values in 90–100 image regions after masking-out SExtractor-detected sources. Typical mean sky values were 0.4 counts pixel^{−1} in FUV and 4 counts pixel^{−1} in NUV. We computed the mean sky value uncertainty in each image from the distribution of sky values in the small regions assuming a Gaussian distribution. These were typically 0.06 counts pixel^{−1} in FUV and 0.2 counts pixel^{−1} in NUV. We took the low-level flat-fielding errors to be negligible, so these measurements form the basis of the errors plotted in the individual FUV and NUV radial surface brightness profiles of Fig. 1. At large radii, the galaxies are faint and only a few counts are detected in each pixel. Accurately determining the level of the sky, even if faint, is thus the main factor limiting the depth of the images and our surface photometry. The UV radial profiles in Fig. 1 are truncated when the uncertainty in the surface brightness exceeds 0.3 mag arcsec^{−2}. The *V*-band profiles are truncated at the same radial distance where their corresponding NUV profiles end.

We convolved the FUV and optical data to the spatial resolution of the NUV observations before any analysis, to avoid spurious colour gradients in the inner parts. We carried out surface photometry by measuring the surface brightness along elliptical annuli, using the iterative method described by Jedrzejewski (1987) and

implemented in the ELLIPSE task within the STSDAS ISOPHOTE package in Image Reduction and Analysis Facility (IRAF). The centre of the isophotes was fixed to the centre of the light distribution, and the position angle, ellipticity (ϵ) and surface brightness (μ) were fitted as a function of the radius. Ellipses were fitted to the *V*-band images only, which have a far superior signal-to-noise ratio (S/N) at all radii compared with the UV images, and were then imposed on the UV images so that meaningful colours could be derived. Galactic extinction was corrected with $R_V = 3.1$ (Cardelli, Clayton & Mathis 1989), $A_{FUV} = 8.376 \times E(B - V)$ and $A_{NUV} = 8.741 \times E(B - V)$ (Wyder et al. 2005) using reddening maps from Schlegel, Finkbeiner & Davis (1998). We present the *GALEX* FUV, NUV and MDM F555W images as well as the surface brightness, colour and ellipticity radial profiles in Fig. 1. The vertical dashed lines in Fig. 1 show the effective radii determined by the SAURON survey at *I* band (see Paper VI). For reference, a UV spectral slope of zero in the $\lambda - F_\lambda$ plane is roughly the same as $FUV - NUV = 1.0$ in AB magnitudes (shown as a horizontal line in the FUV–NUV colour profiles of Fig. 1). We give comments on notable features in the profiles of individual galaxies in Appendix A.

From the radial surface brightness profiles, we derived total apparent magnitudes by extrapolating the growth curves to infinity. For each profile, we first calculated the integrated apparent magnitude $[m(R)]$ within elliptical isophotes up to the radius where the uncertainty in the surface brightness reaches 0.3 mag arcsec^{−2}. We then computed the slope of the growth curve $[S \equiv dm(R)/dR]$ within that region, plotted it as a function of the integrated magnitude $m(R)$, and fitted a straight line to the outer parts. The value of m where the extrapolated value of S is 0 was taken as the total apparent magnitude. Selecting a suitable radial range for growth curve fitting was occasionally difficult in FUV because of the low S/N in the outer parts, but the errors quoted correspond to the uncertainties associated with the fit to the (slope of the) growth curve only, so they are probably slightly underestimated. This technique is described in more detail in Cair3s et al. (2001) and Gil de Paz et al. (2007). The effective radius (R_e ; semimajor axis of the elliptical aperture containing half the light), the mean surface brightness within R_e ($\langle \mu \rangle_e$) and their uncertainties were also calculated from the growth curve. Having said that, the uncertainty in R_e arising from different measurement methods is always larger than formal uncertainties, and a comparison of different measurements shows that it is typically ≈ 20 per cent, which we adopt (see Falc3n-Barroso et al., in preparation).

In Fig. 2, we show the comparison of our magnitudes and effective radii in both FUV and NUV with those of Gil de Paz et al. (2007). The agreement is generally good. The mean of the absolute values of the total apparent magnitude differences is ≈ 0.16 mag in both FUV and NUV, as expected generally somewhat larger than our formal uncertainties. For the effective radii, the mean absolute difference is ≈ 9 per cent in FUV and ≈ 15 per cent in NUV, so generally within our errors. The integrated UV and other properties are listed in Table 1.

3 DISCUSSION

3.1 UV colour–magnitude relations

The simplest new data product obtained from our *GALEX* observations is the integrated UV apparent magnitudes, from which we derive total absolute magnitudes using the distances listed in Table 1. CMRs are widely used to study the star formation history (SFH) of early-type galaxies. The NUV CMR (see e.g. Yi et al.

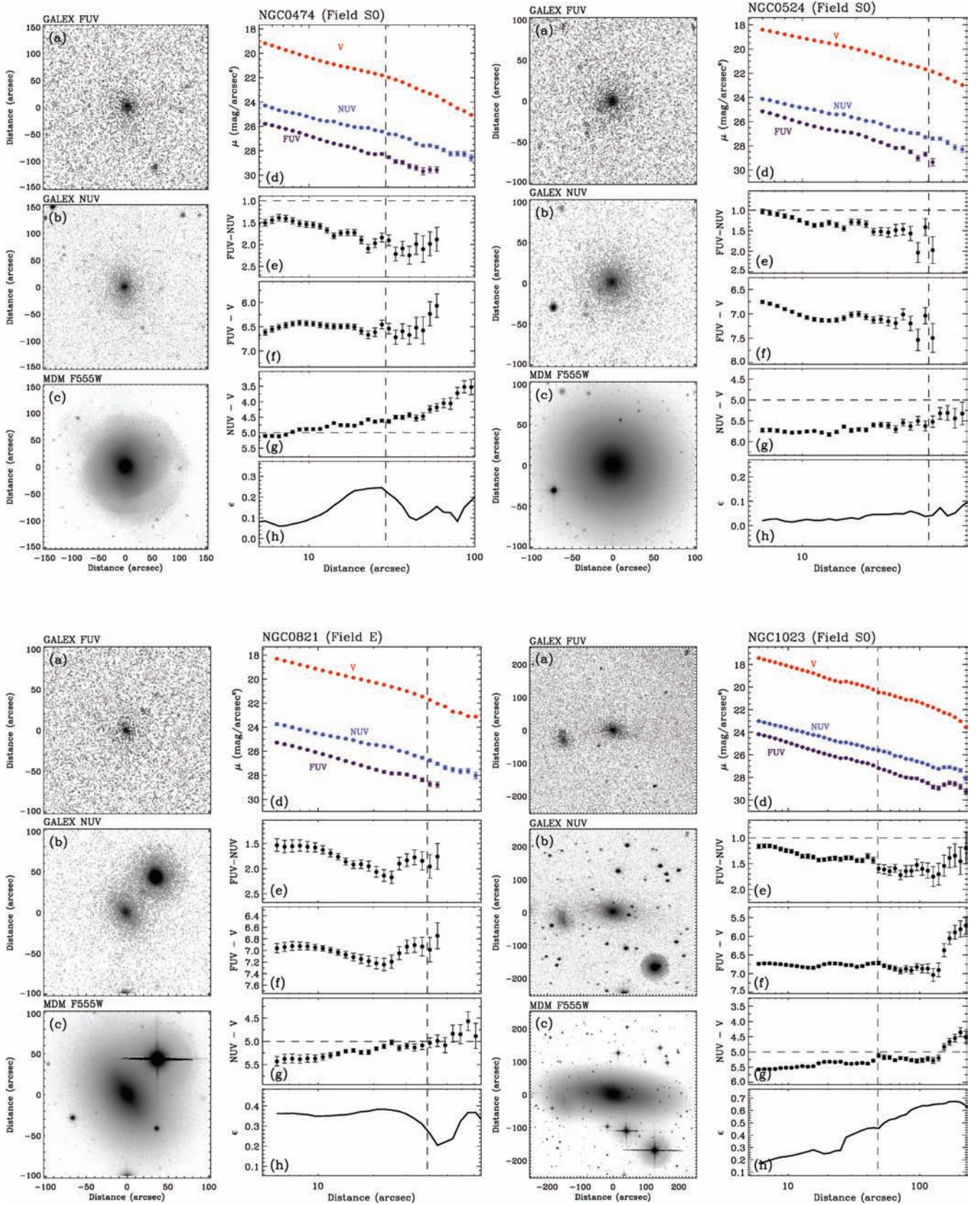


Figure 1. *GALEX* UV and MDM optical images and surface photometry of 34 early-type galaxies in the SAURON sample. (a)–(c) FUV, NUV and F555W images. (d) V, NUV and FUV radial surface brightness profiles. (e)–(g) FUV–NUV, FUV–V and NUV–V radial colour profiles. (h) Ellipticity radial profile derived from the optical image. The two horizontal lines show FUV–NUV = 1.0 and NUV–V = 5.0, respectively. The vertical lines show the *I*-band effective radius from the SAURON survey (Paper VI).

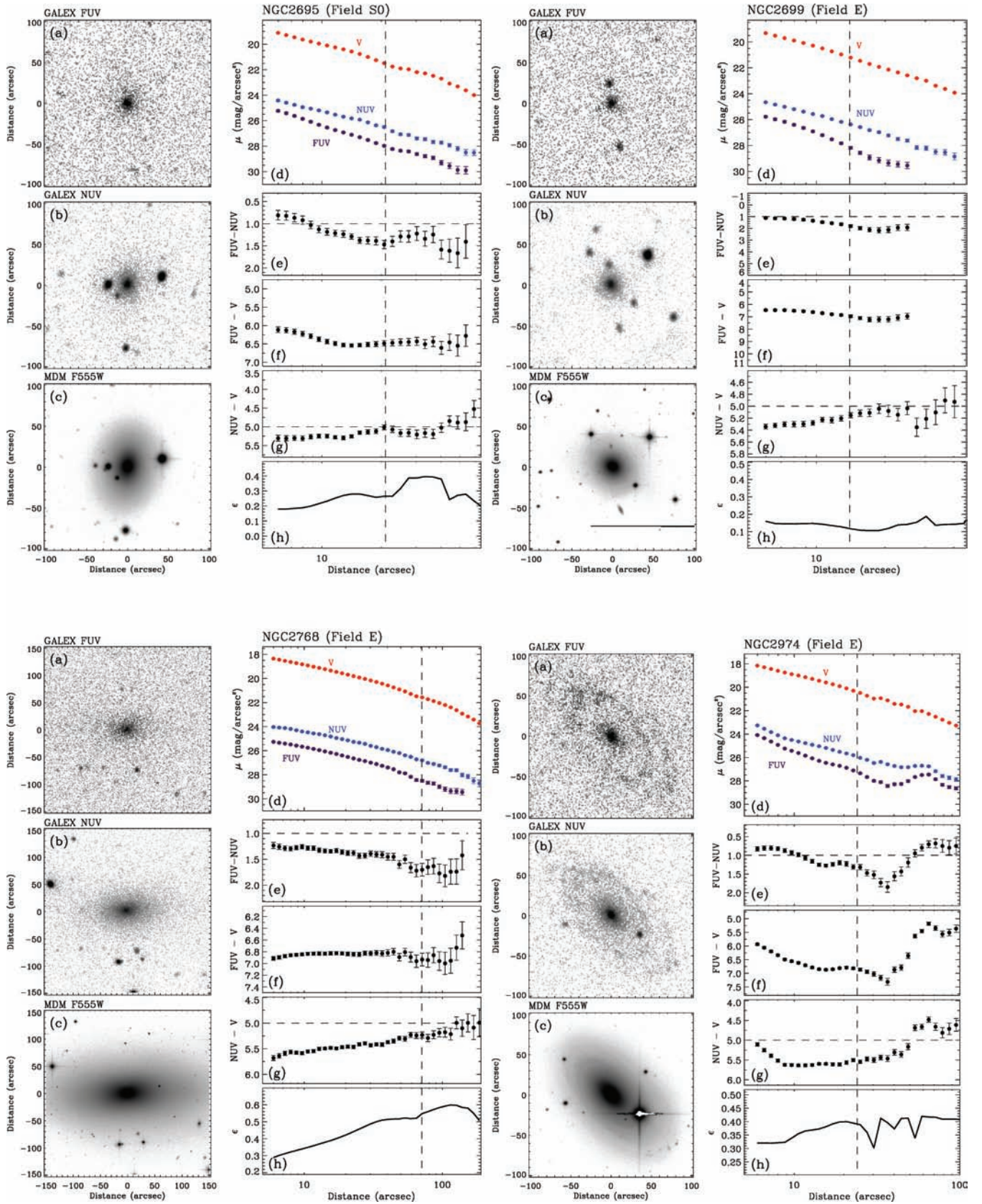


Figure 1 – continued

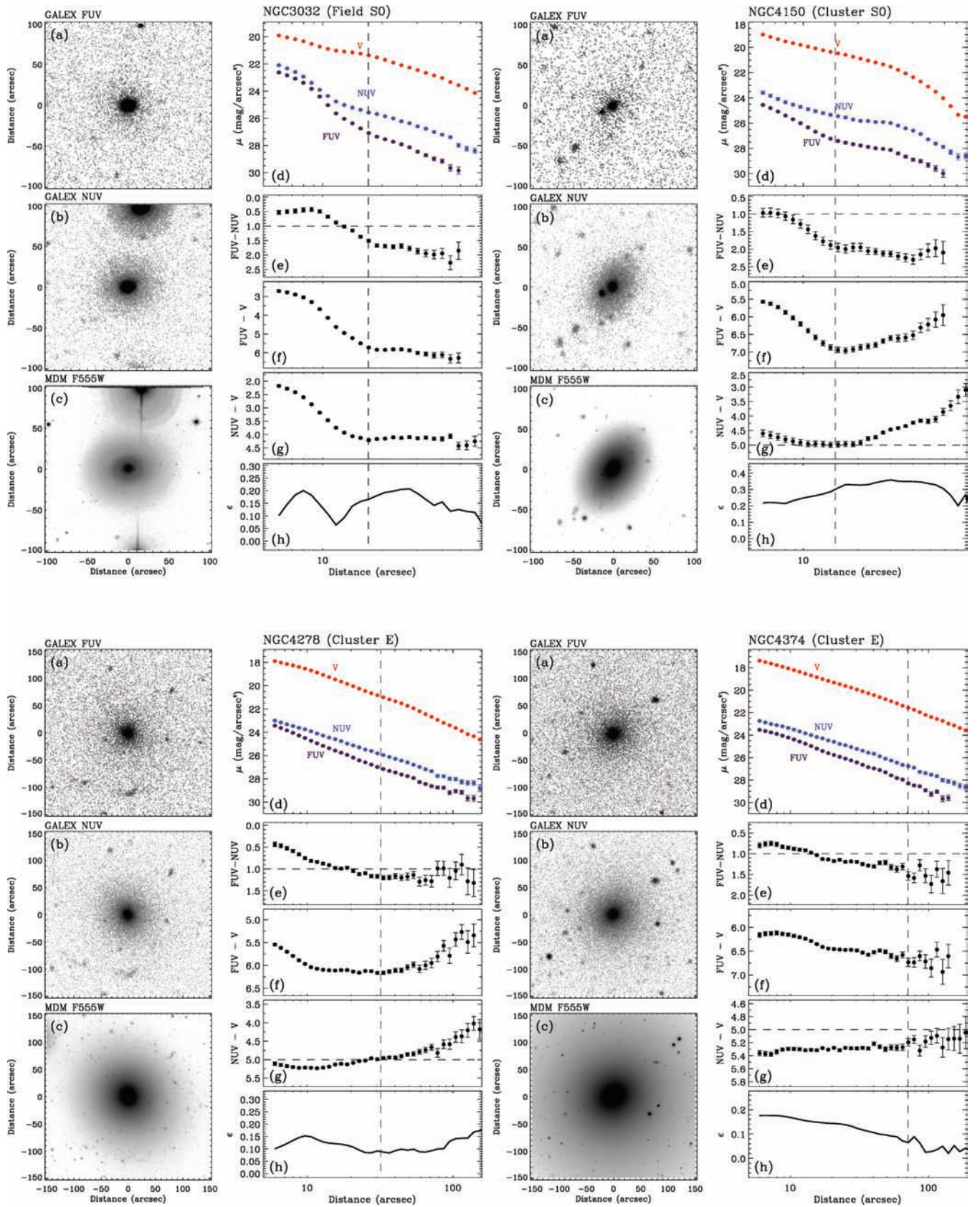


Figure 1 – continued

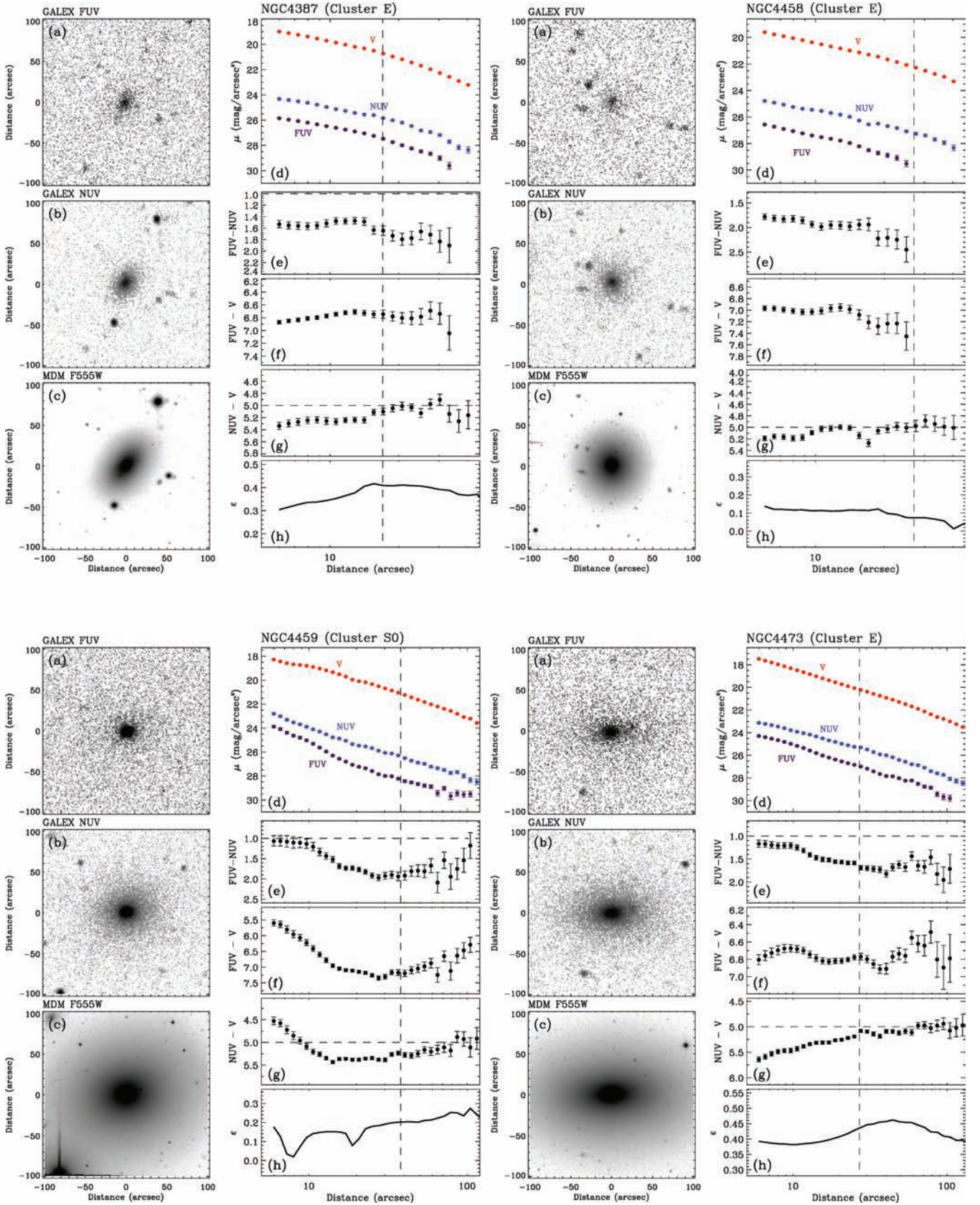


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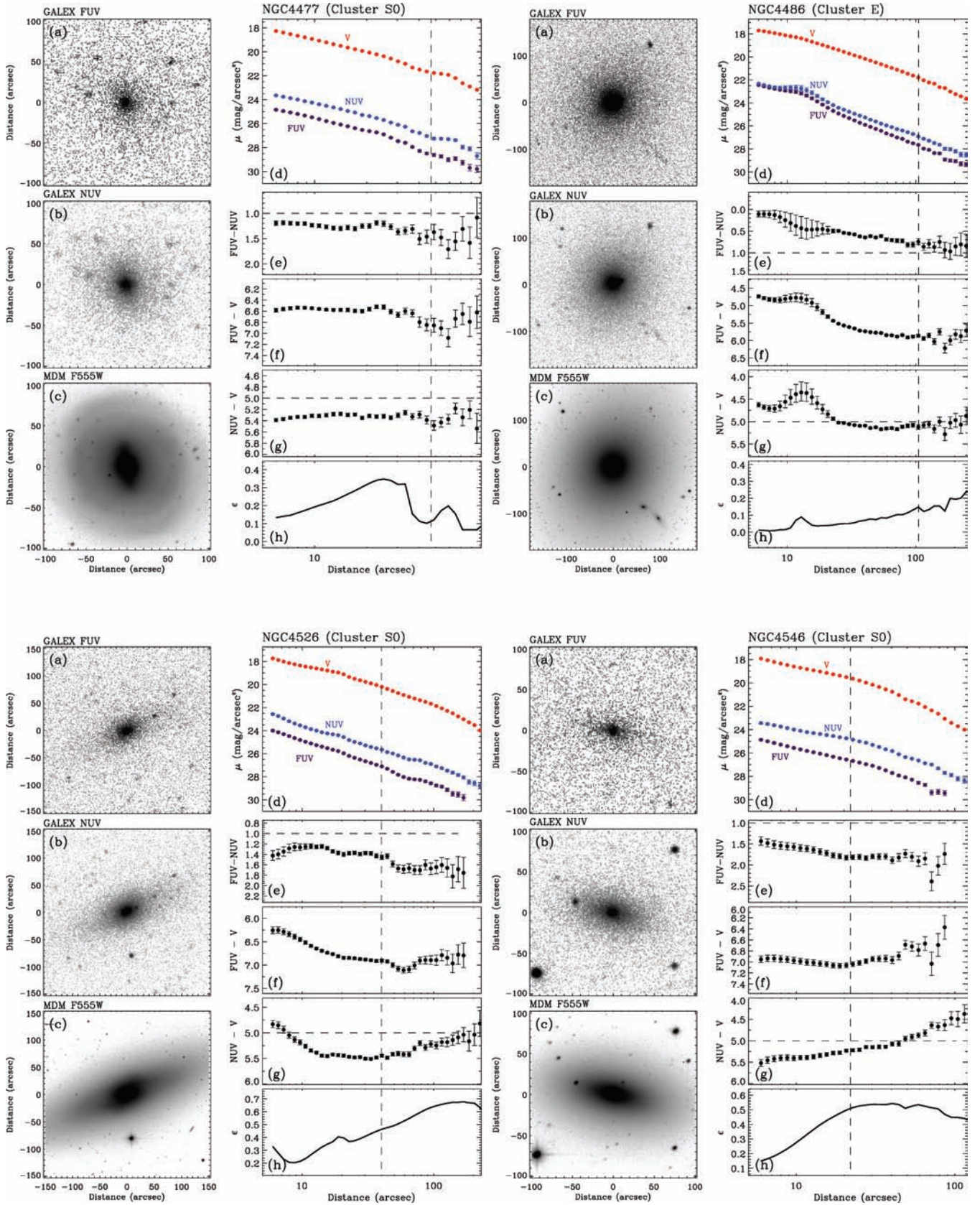


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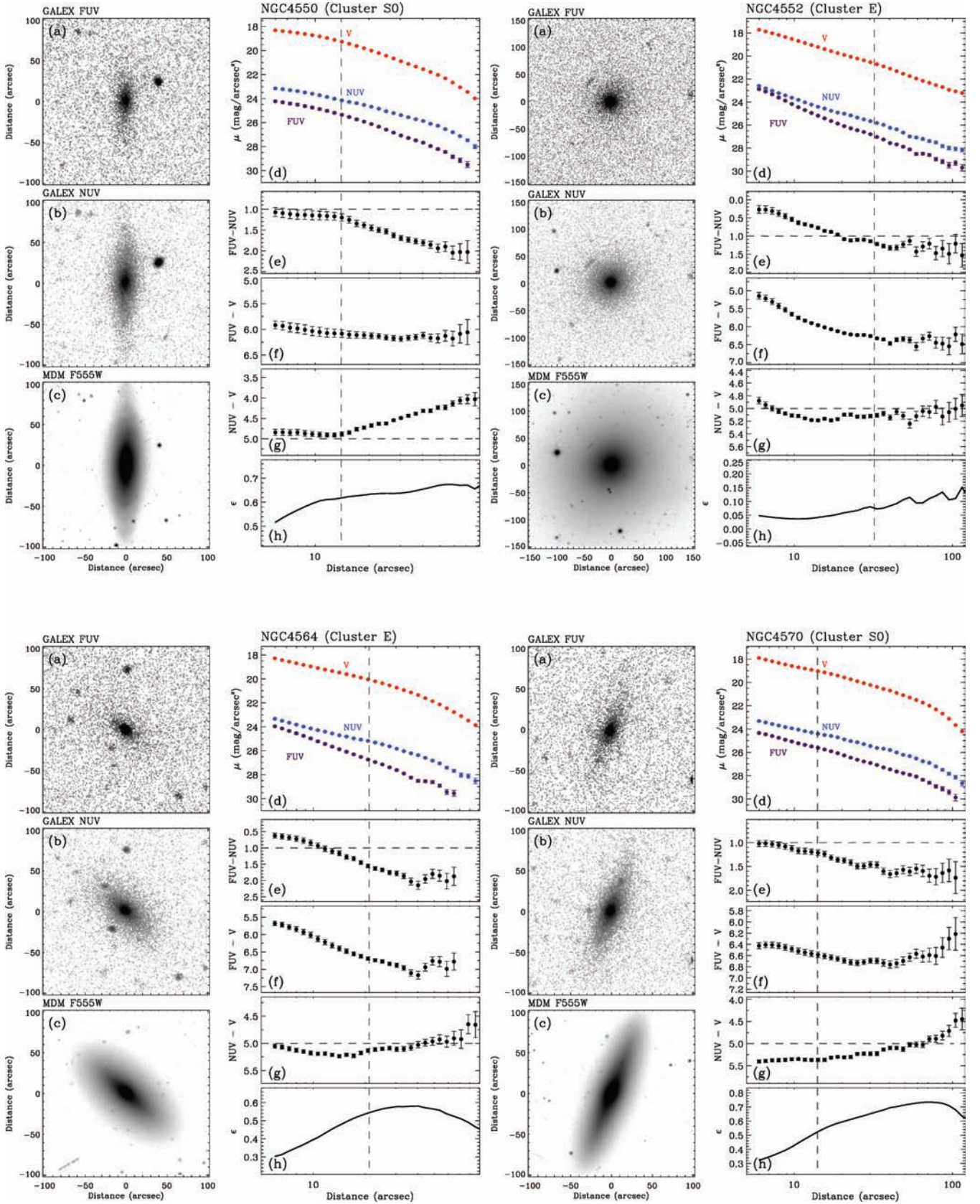


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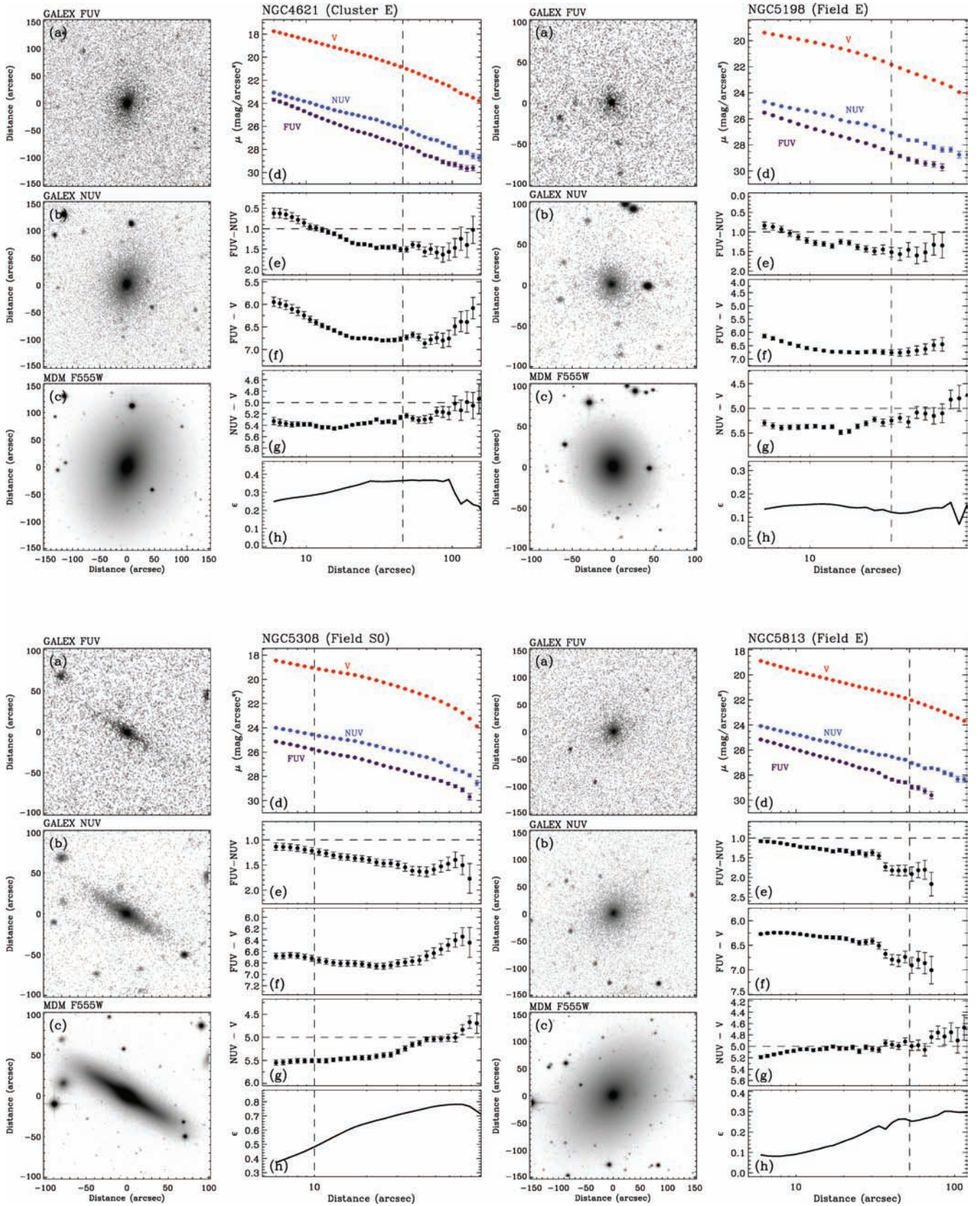


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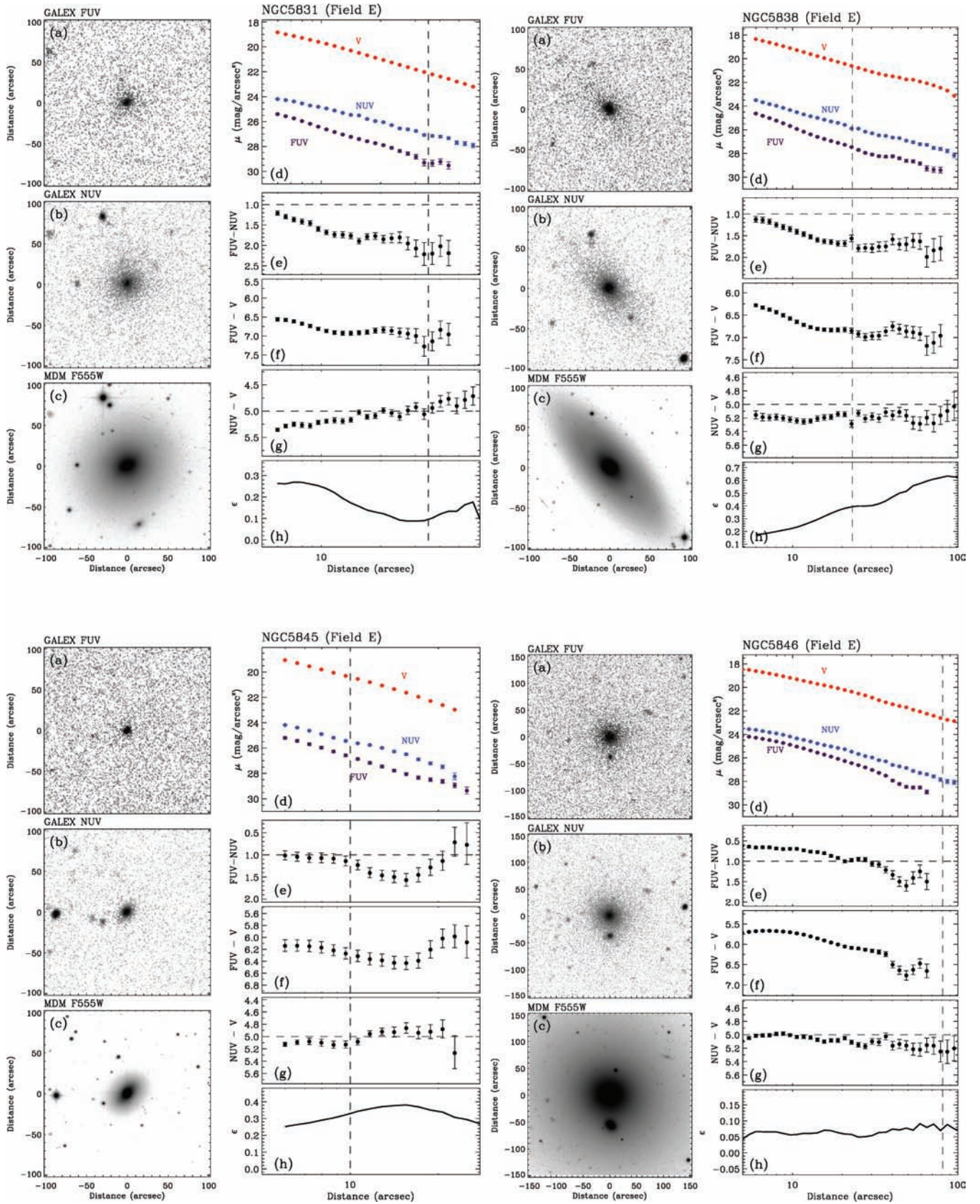


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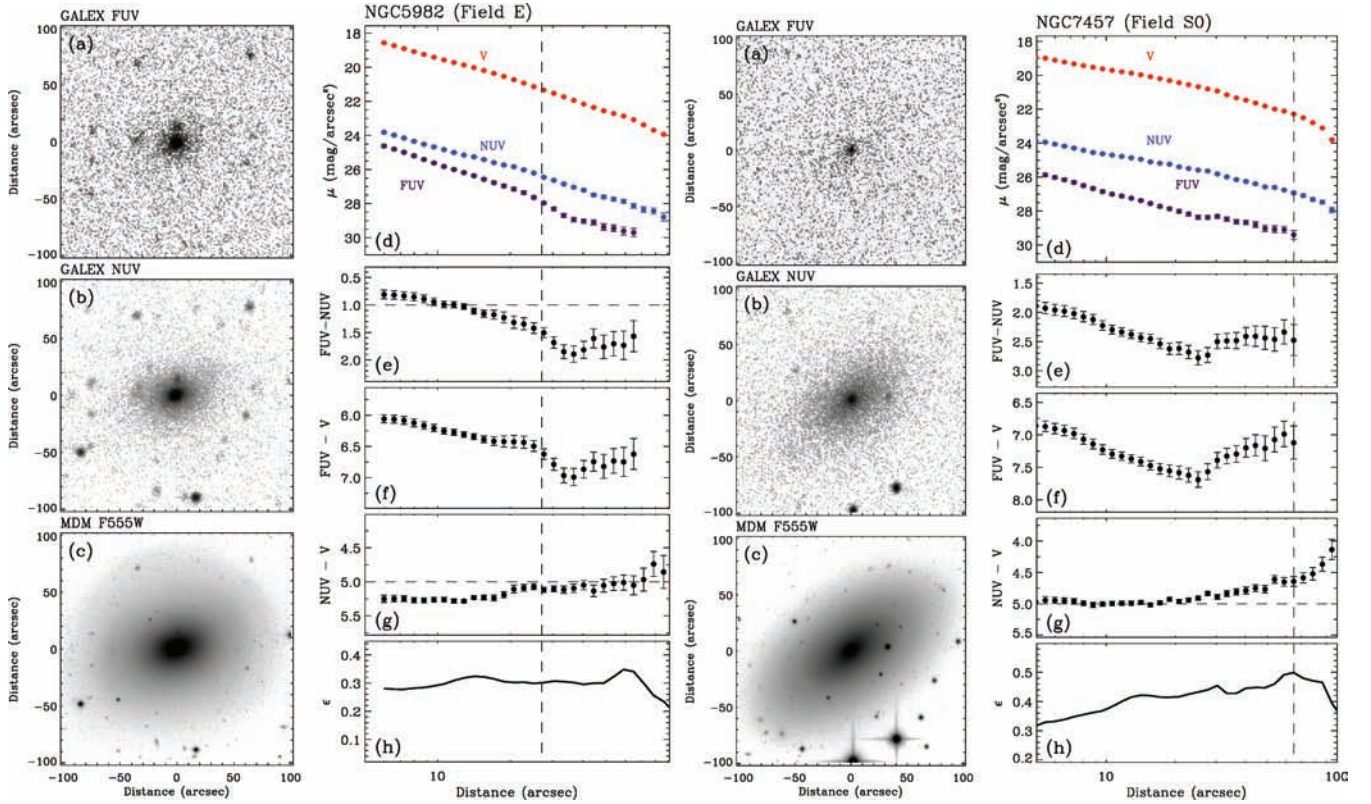


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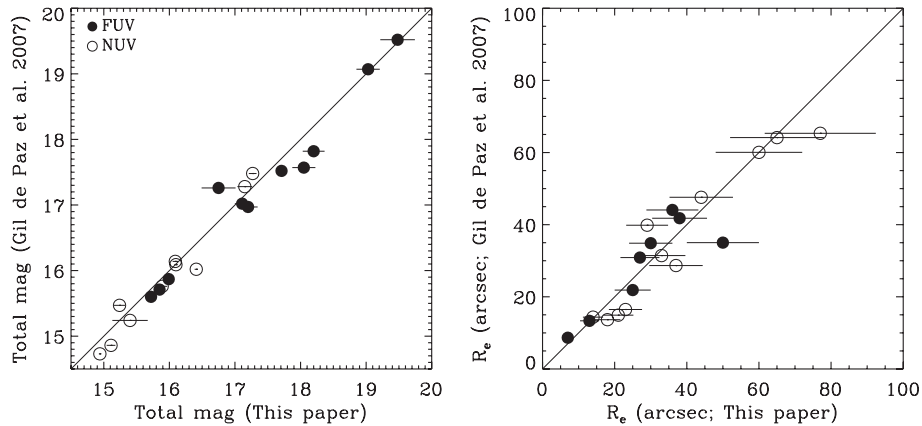


Figure 2. Comparison of our magnitudes and effective radii in the FUV and NUV passbands with those of Gil de Paz et al. (2007).

2005; Kaviraj et al. 2007) is a particularly good tool for tracking RSF, owing to its high sensitivity to the presence of young stellar populations. Fig. 3 shows the UV–optical CMRs. It also shows the red sequence NUV– r fit of Yi et al. (2005), after correcting for the sky measurement offset from the *GALEX* pipeline (see Gil de Paz et al. 2007). The scatter in the CMRs of Fig. 3 is significant and there are clearly outliers about the red sequence relation. We henceforth attempt to quantify the effects of RSF on those relations.

Most UV photons are emitted by stars younger than a few hundred megayears, so they are useful to trace residual star formation, but FUV flux can also be generated by evolved hot helium-burning stars

(the so-called UV-upturn phenomenon; see e.g. Yi, Demarque & Oemler 1997; O’Connell 1999 for reviews). Star formation studies thus usually use the NUV passband instead (Yi et al. 2005), although it is nevertheless important to know how much of the NUV flux can originate from the UV-upturn phenomenon. NGC 4552 is a famous UV-upturn galaxy; its FUV– V colour is one of the bluest among nearby elliptical galaxies with no sign of RSF. We measure an NUV– V colour of 5.15 ± 0.03 mag for NGC 4552, so considering the measurement error we adopt a very conservative 5σ upper limit of NUV– $V = 5.0$ to pick out galaxies that have experienced RSF. This empirical demarcation for identifying RSF galaxies is shown as a horizontal dashed line in the NUV– V colour profiles of Fig. 1 and

Table 1. Integrated properties.

Galaxy	$E(B - V)$	D (Mpc)	σ_e (km s ⁻¹)	R_e		$\langle\mu\rangle_e$		Total apparent magnitude		UV type
				FUV (arcsec)	NUV	FUV (mag)	NUV	FUV (mag)	NUV (mag)	
NGC 0474	0.034	26.1	150	74	67	28.96 ± 0.49	26.96 ± 0.31	17.78 ± 0.37	15.82 ± 0.28	RSF
NGC 0524	0.083	24.0	235	17	39	25.90 ± 0.30	25.82 ± 0.26	17.70 ± 0.13	15.88 ± 0.15	
NGC 0821	0.110	24.1	189	27	33	27.21 ± 0.50	25.84 ± 0.27	18.10 ± 0.36	15.77 ± 0.12	A/Z
NGC 1023	0.061	11.4	182	44	70	26.38 ± 0.30	25.76 ± 0.25	16.16 ± 0.15	14.53 ± 0.12	RSF
NGC 2695	0.018	32.4	188	13	27	25.89 ± 0.28	25.96 ± 0.24	18.32 ± 0.11	16.81 ± 0.11	UVX, A/Z
NGC 2699	0.020	26.9	124	9	27	25.93 ± 0.29	26.32 ± 0.26	19.06 ± 0.14	17.27 ± 0.16	A/Z
NGC 2768	0.044	22.4	216	50	77	27.86 ± 0.38	26.82 ± 0.31	16.75 ± 0.26	15.35 ± 0.27	
NGC 2974	0.054	21.5	233	11	29	24.91 ± 0.27	25.43 ± 0.23	17.71 ± 0.09	16.10 ± 0.09	RSF
NGC 3032	0.017	22.0	90	2	3	20.83 ± 0.25	20.43 ± 0.20	16.94 ± 0.05	15.78 ± 0.01	RSF
NGC 4150	0.018	13.7	77	5	18	23.46 ± 0.25	24.32 ± 0.21	18.20 ± 0.17	16.10 ± 0.03	RSF
NGC 4278	0.029	16.1	231	27	44	23.63 ± 0.25	24.49 ± 0.21	16.41 ± 0.08	15.21 ± 0.07	UVX, A/Z
NGC 4374	0.040	18.5	278	38	65	25.63 ± 0.27	25.42 ± 0.23	15.72 ± 0.09	14.35 ± 0.08	UVX
NGC 4387	0.033	18.0	98	12	15	26.33 ± 0.31	25.08 ± 0.22	19.03 ± 0.18	17.27 ± 0.06	
NGC 4458	0.024	16.4	85	13	23	27.08 ± 0.41	25.95 ± 0.25	19.40 ± 0.26	16.84 ± 0.11	
NGC 4459	0.046	16.1	168	13	20	24.63 ± 0.26	24.13 ± 0.21	17.02 ± 0.07	15.49 ± 0.03	RSF
NGC 4473	0.028	15.3	192	25	47	26.09 ± 0.29	25.60 ± 0.23	17.11 ± 0.11	15.24 ± 0.08	
NGC 4477	0.032	16.7	162	30	33	27.36 ± 0.40	26.04 ± 0.25	17.24 ± 0.15	15.48 ± 0.08	
NGC 4486	0.022	17.2	298	36	51	24.29 ± 0.26	24.46 ± 0.21	14.50 ± 0.06	13.92 ± 0.03	UVX
NGC 4526	0.022	16.9	222	25	26	25.60 ± 0.27	24.24 ± 0.21	16.59 ± 0.09	15.22 ± 0.03	RSF
NGC 4546	0.034	14.1	194	23	54	26.33 ± 0.31	25.85 ± 0.24	17.56 ± 0.13	15.41 ± 0.13	A/Z
NGC 4550	0.039	15.5	110	7	14	24.30 ± 0.26	24.05 ± 0.21	18.05 ± 0.18	16.41 ± 0.03	RSF
NGC 4552	0.041	15.9	252	13	19	23.39 ± 0.25	23.59 ± 0.21	15.85 ± 0.06	14.94 ± 0.02	UVX
NGC 4564	0.035	15.9	155	5	21	23.15 ± 0.25	24.65 ± 0.21	17.71 ± 0.06	16.09 ± 0.04	UVX, A/Z
NGC 4570	0.022	17.1	173	13	26	25.00 ± 0.26	25.00 ± 0.22	17.52 ± 0.08	15.95 ± 0.05	A/Z
NGC 4621	0.033	14.9	211	11	35	23.83 ± 0.26	24.93 ± 0.22	16.65 ± 0.06	15.25 ± 0.05	UVX
NGC 5198	0.023	38.4	179	7	18	24.92 ± 0.27	25.47 ± 0.23	18.86 ± 0.11	17.23 ± 0.08	UVX, A/Z
NGC 5308	0.018	29.4	208	12	25	25.75 ± 0.28	25.66 ± 0.23	18.36 ± 0.11	16.81 ± 0.09	A/Z
NGC 5813	0.057	32.2	230	40	80	27.20 ± 0.35	26.82 ± 0.32	17.19 ± 0.24	15.50 ± 0.34	A/Z
NGC 5831	0.059	27.2	151	15	55	26.23 ± 0.30	26.70 ± 0.31	18.36 ± 0.15	16.20 ± 0.32	A/Z
NGC 5838	0.053	24.7	240	10	21	24.81 ± 0.27	24.73 ± 0.22	17.77 ± 0.08	16.16 ± 0.05	
NGC 5845	0.053	25.9	239	6	6	24.59 ± 0.27	23.59 ± 0.21	18.78 ± 0.12	17.70 ± 0.02	A/Z
NGC 5846	0.055	24.9	238	22	32	25.37 ± 0.27	25.15 ± 0.23	16.66 ± 0.09	15.51 ± 0.07	UVX
NGC 5982	0.018	51.7	229	12	29	25.19 ± 0.27	25.58 ± 0.23	17.76 ± 0.08	16.25 ± 0.08	UVX, A/Z
NGC 7457	0.052	13.2	78	37	43	26.35 ± 0.30	25.89 ± 0.25	18.39 ± 0.31	15.83 ± 0.13	RSF

Columns: (1) Galaxy identifier. (2) Galactic extinction in the NED from Schlegel et al. (1998). (3) Distance from Falcón-Barroso et al. (in preparation). (4) Velocity dispersion of the luminosity-weighted spectrum within a circular aperture of R_e from Paper IX. The uncertainty is taken as 5 per cent. (5)–(6) Effective radius in each band. The uncertainties are taken as 20 per cent. (7)–(8) Mean effective surface brightness in each band and error. (9)–(10) Total apparent magnitude in each band and error. (11) UV–optical radial colour profiles classification. RSF: recent star formation. UVX: UV upturn. A/Z: large-scale age and/or metallicity gradient.

the NUV–V CMR of Fig. 3. We henceforth assume that the relative FUV strength of NGC 4552 is the maximum that can be generated by purely old-star populations. This is an ad hoc assumption; however, it is frequently made and is empirically grounded.

Using this empirical criterion, we have identified the galaxies that are likely to have experienced a recent episode of star formation in Fig. 3 onwards (filled symbols). We label galaxies with $\text{NUV} - V \geq 5.0$ as quiescent (QST) early-type galaxies (open symbols). Furthermore, stars and concentric circles indicate Balmer absorption line strengths $H\beta \gtrsim 2$ Å and field galaxies, respectively. The $H\beta$ indices are computed within one effective radius in the optical and come from the SAURON data published in Paper VI. Supporting our empirical threshold, four out of five RSF candidate galaxies (NGC 3032, 4150, 4550 and 7457) also show enhanced $H\beta$ line strengths, a widely used post-starburst signature. The only exception (NGC 474) is noted for its shell structures located in the outer regions, at $R \approx 60$ arcsec, and as shown in Fig. 1 the total NUV flux is dominated by these outer regions, whereas the Balmer line was measured within R_e in the optical.

A more in-depth look at Fig. 3, with likely RSF galaxies identified, reveals that a non-negligible fraction of the scatter in the UV CMRs (and the departure from the Yi et al. 2005 red sequence relation) is due to galaxies with RSF. For example, the scatter in the $(\text{NUV} - V) - M_V$ CMR increases by 67 per cent between least-squares fits excluding and including RSF candidates. The overall fraction of galaxies with RSF is 15 per cent (5/34). This may appear to also agree with the result of Yi et al. (2005), but a direct comparison is not sensible because the sampling strategies are radically different. It should also be noted that no clear dependence of $\text{FUV} - V$ or $\text{FUV} - \text{NUV}$ on M_V is visible in our sample. Finally, environment does not appear to play a significant role in star formation, as the RSF galaxies selected contain both field and cluster galaxies. We recall however that the densest environment probed by the SAURON survey is that of the Virgo cluster of galaxies (see Paper II), and that our sample is admittedly too small to robustly investigate environmental effects.

We also plot the colour– σ_e (integrated stellar velocity dispersion within R_e) and colour–mass (virial mass estimate) relations

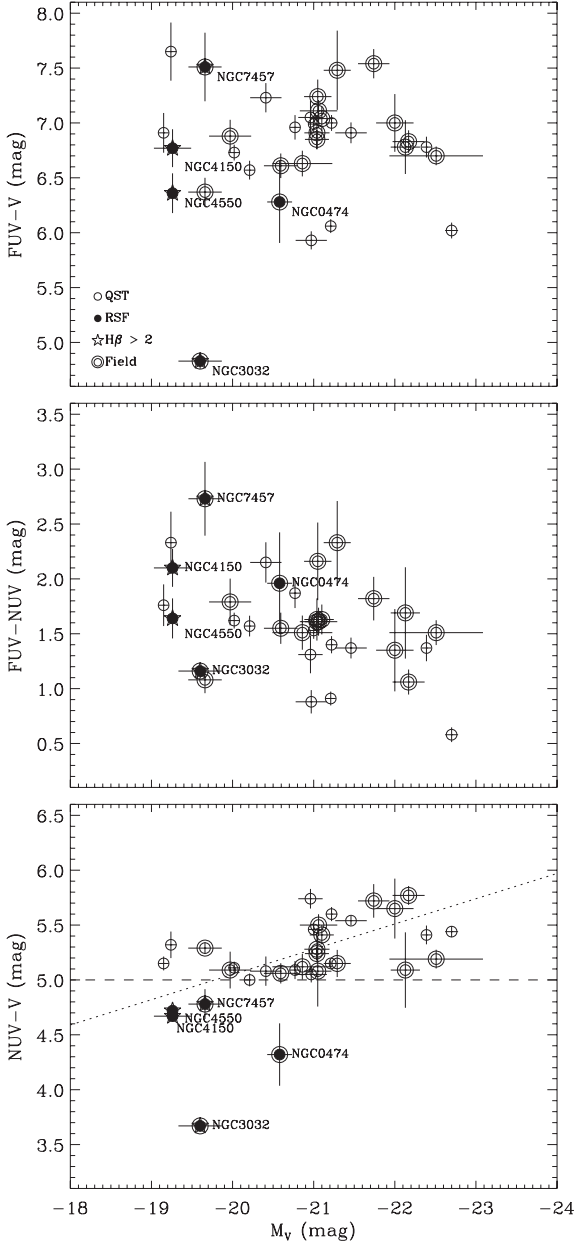


Figure 3. UV CMRs of our 34 SAURON early-type galaxies. The FUV–V (top panel), FUV–NUV (middle panel) and NUV–V (bottom panel) colours are shown as a function of the absolute magnitude in V. The dashed line indicates the NUV–V = 5.0 cut-off for RSF. Using this standard, we divided the sample into quiescent early-type galaxies (QST; open symbols) and RSF (filled symbols) galaxies. Stars denote galaxies with $H\beta \gtrsim 2 \text{ \AA}$. Field galaxies are represented by two concentric circles. The red-sequence fit from Yi et al. (2005) in NUV–r is shown as a dotted line in the NUV–V CMR (with offset applied; see the text).

in Fig. 4. The dynamical mass estimates are of the form $M_{\text{vir}} = (5.0 \pm 0.1) \times R_e \sigma_e^2 / G$ (Paper IV). RSF galaxies generally have lower velocity dispersions and correspondingly smaller virial masses than the majority of red early-type galaxies, and there is no RSF early-type galaxy with a velocity dispersion $\sigma_e > 200 \text{ km s}^{-1}$ (see also Schawinski et al. 2006). This implies that only low-mass early-type galaxies have formed stars recently, confirming the results of Schawinski et al. (2007). In addition, the more massive (high-velocity dispersion) early-type galaxies tend to be

bluer in both FUV–V and FUV–NUV, confirming earlier studies (e.g. Burstein et al. 1988; Donas et al. 2007).

Interestingly, using *Spitzer* Infrared Array Camera imaging and Infrared Spectrograph spectroscopy, Shapiro et al. (in preparation) surveyed the mechanisms driving star formation in the SAURON early-type galaxies, and found that star formation happens exclusively in fast-rotating systems (see Paper IX). All of our RSF galaxies are also classified as star-forming galaxies in their survey, except again for NGC 474, and all are fast-rotating except NGC 4550, the peculiar galaxy with co-spatial counter-rotating discs (extensively discussed in Paper IX).

Our star formation interpretation is also consistent with the results of surveys of the molecular gas emission in early-type galaxies by Combes, Young & Bureau (2007), Crocker et al. (2008) and Young, Bureau & Cappellari (2008). CO emission is detected in three of the RSF candidates (NGC 3032, 4150 and 4550). In the case of NGC 7457, a molecular gas mass of $4.4 \times 10^6 M_\odot$ was reported by Welch & Sage (2003), but Combes et al. (2007) later failed to detect it. NGC 474 was also not detected by Combes et al. (2007), but the 24 arcsec primary beam did not encompass the shells and strong NUV emission. We summarize the CO detection rates of our sample galaxies in Table 2, where we have divided the sample into three categories: RSF; partial RSF and quiescent galaxies. Again, RSF are galaxies selected by the empirical criterion $\text{NUV} - V < 5.0$. Partial RSF galaxies do not satisfy the overall blue colour criterion in integrated NUV–V, but they show significant blue regions in their radial profiles. Quiescent galaxies have $\text{NUV} - V \geq 5.0$ at essentially all radii. The CO detection rate decreases monotonically from RSF to quiescent galaxies, thus supporting our RSF identifications and suggesting, unsurprisingly, that CO-rich galaxies tend to have had more RSF.

3.2 UV fundamental planes

Early-type galaxies follow a correlation between total luminosity and central velocity dispersion of the form $L \propto \sigma_e^q$, where $q \approx 4$ with considerable scatter. We present this Faber–Jackson relation (Faber & Jackson 1976) for our sample galaxies in Fig. 5 for both FUV and NUV bands, using the same symbols as in Fig. 3 and with a line of slope 4 overlaid. It is striking that the galaxies previously identified as RSF are among the greatest outliers in both relations, where they are responsible for much of the total scatter. When RSF early-type galaxies are included, the scatter increases by 65 and 69 per cent in the FUV and NUV, respectively, compared to that when only quiescent galaxies are considered.¹ This is consistent with the conclusions from our analysis of the CMRs in the previous section.

We now consider the FP (Djorgovski & Davis 1987; Dressler et al. 1987), a relation between the photometric effective radius (R_e), the mean surface brightness within R_e ($\langle \mu \rangle_e$) and the effective stellar velocity dispersion (σ_e) of the form

$$\log(R_e) = a \log(\sigma_e) + b \log(\langle \mu \rangle_e) + c, \quad (1)$$

where a , b and c are constant coefficients derived by minimizing the residuals from the plane, and $a = 2.0$ and $b = 0.4$ are expected from the virial theorem if the galaxies form a homologous family with a constant mass-to-light ratio. Just as for Faber–Jackson, our velocity dispersions are measured in the optical (roughly V band),

¹ For reference, simple linear fits to the whole sample yield the slopes of $3.59^{+0.64}_{-0.47}$ (FUV) and $3.09^{+0.73}_{-0.50}$ (NUV), whereas for the quiescent subsample we find $4.12^{+0.60}_{-0.46}$ (FUV) and $3.63^{+0.74}_{-0.52}$ (NUV).

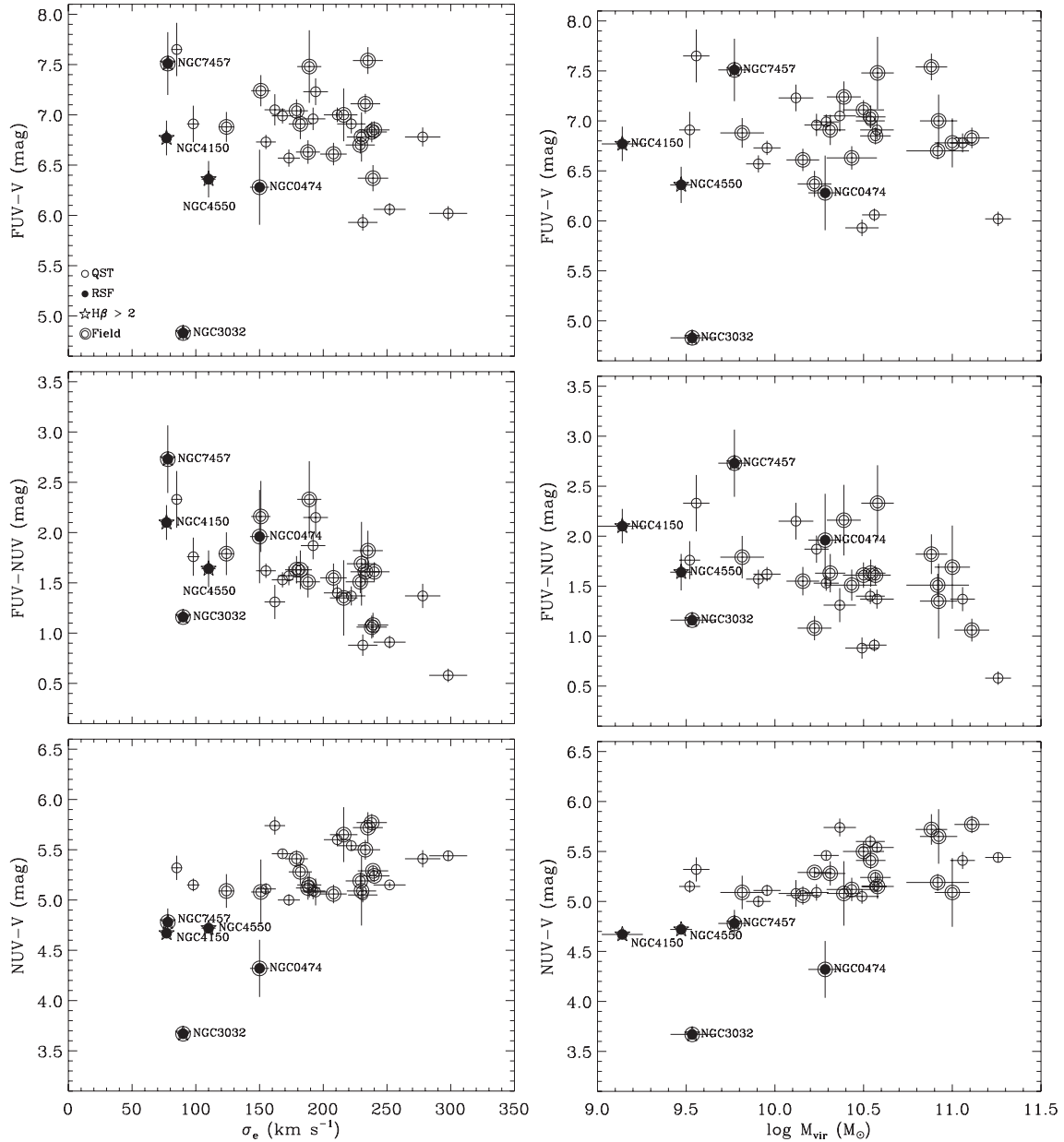


Figure 4. Colour–stellar velocity dispersion and colour–mass relations of our 34 SAURON early-type galaxies. Symbols are the same as in Fig. 3.

Table 2. CO detection rates in RSF and quiescent early-type galaxies.

	RSF	Partial RSF	QST
Number of sample galaxies	5	4	25
Galaxies detected in CO	3	2	4
Rate (per cent)	60 ± 22	50 ± 25	16 ± 7

Note. Errors assume a binomial distribution.

while our photometric quantities are measured in the UV. The FP of course implicitly relies on all quantities used in the virial theorem being measured for the same stellar population. Our UV FPs are thus somewhat hybrid, and one should keep in mind that this may introduce some biases (because of both differing apertures and wavebands). However, as UV-derived stellar kinematics are

still some time away, this approach is necessary and nevertheless illuminating.

The coefficients a and b measured by Jørgensen et al. (1996) are 1.24 ± 0.07 and 0.33 ± 0.02 in the Gunn r band, respectively. Early investigations suggested that a variation of the coefficient a in different filters would be expected if there were a systematic variation of the mass-to-light ratio of the stellar populations with galaxy luminosity or mass (e.g. Prugniel & Simien 1996). Pahre, De Carvalho & Djorgovski (1998b) indeed reported an increase of the slope with increasing wavelength (from U to K), although Bernardi et al. (2003) found that the FP coefficients were approximately the same in g , r , i and z bands using a much larger but more distant sample of early-type galaxies from SDSS. Meanwhile, Trujillo, Burkert & Bell (2004) suggested that the FP tilt is mostly driven by non-homology and that stellar population effects account for only a small fraction of it. Early type galaxies in field

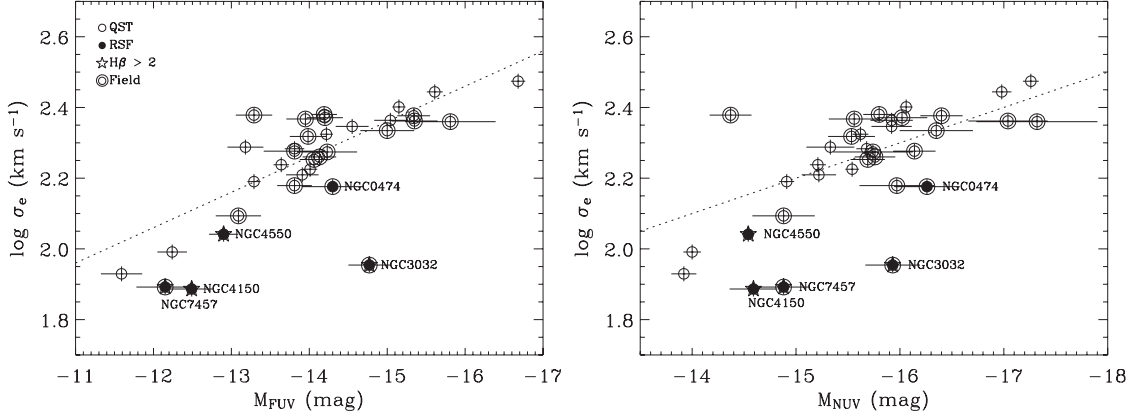


Figure 5. Faber–Jackson relation for our sample of 34 SAURON early-type galaxies in the FUV (top panel) and NUV (bottom panel) passbands. Symbols are the same as in Fig. 3. The original Faber–Jackson relation ($\sigma \propto L^{1/4}$) is shown as a dotted line in each panel.

environments are also generally more diverse and are found to show a greater scatter in the FP (see e.g. De Carvalho & Djorgovski 1992; Zepf & Whitmore 1993). These authors thus suggested that environment plays an important role in the process of galaxy formation and evolution. Recently, new approaches have aimed to elucidate the FP tilt, for example via gravitational lensing (Treu et al. 2006; Bolton et al. 2008) and stellar dynamical modelling (Paper IV). These studies differ from previous ones in that they do not depend on simple virial assumptions, but measure accurate

dynamical masses directly. Using two different approaches, they consistently conclude that the tilt is almost entirely due to a genuine mass-to-light ratio variation and not to non-homology. However, it still remains unclear just how much of this variation can be attributed to a change in the dark matter fraction or to differences in the stellar populations.

For the first time, we present in Figs 6 and 7 the FPs in the FUV and NUV bands, using again the same symbols as in Fig. 3. Table 3 lists the FP coefficients and scatter at both FUV and NUV, obtained

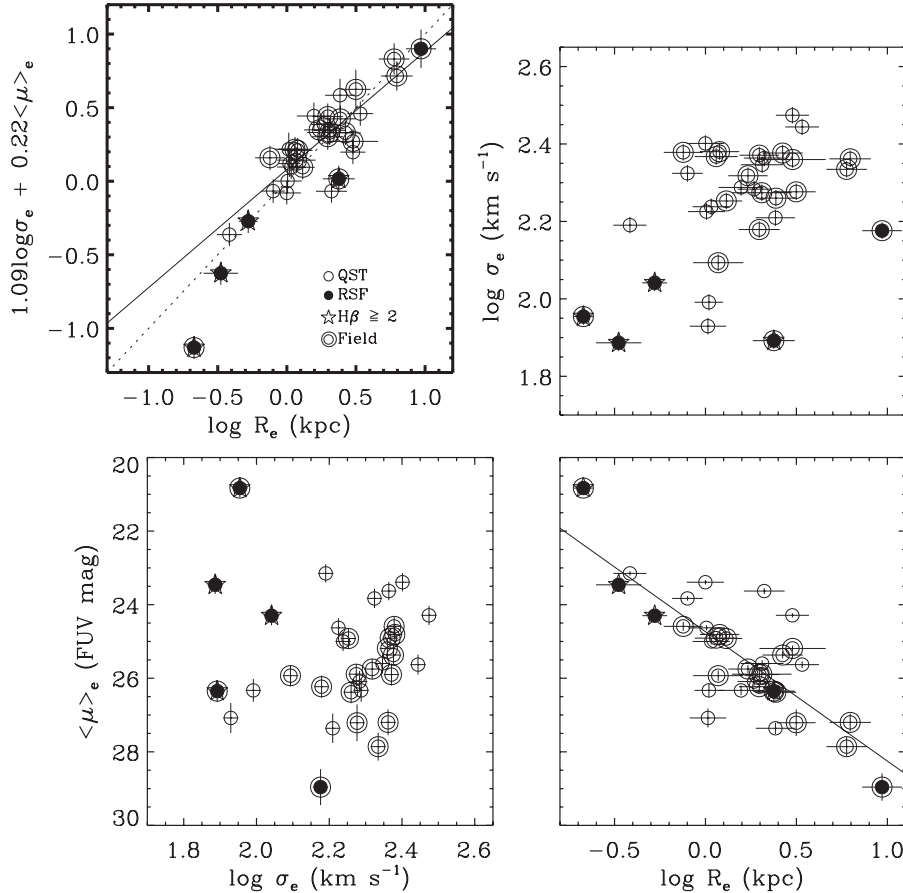
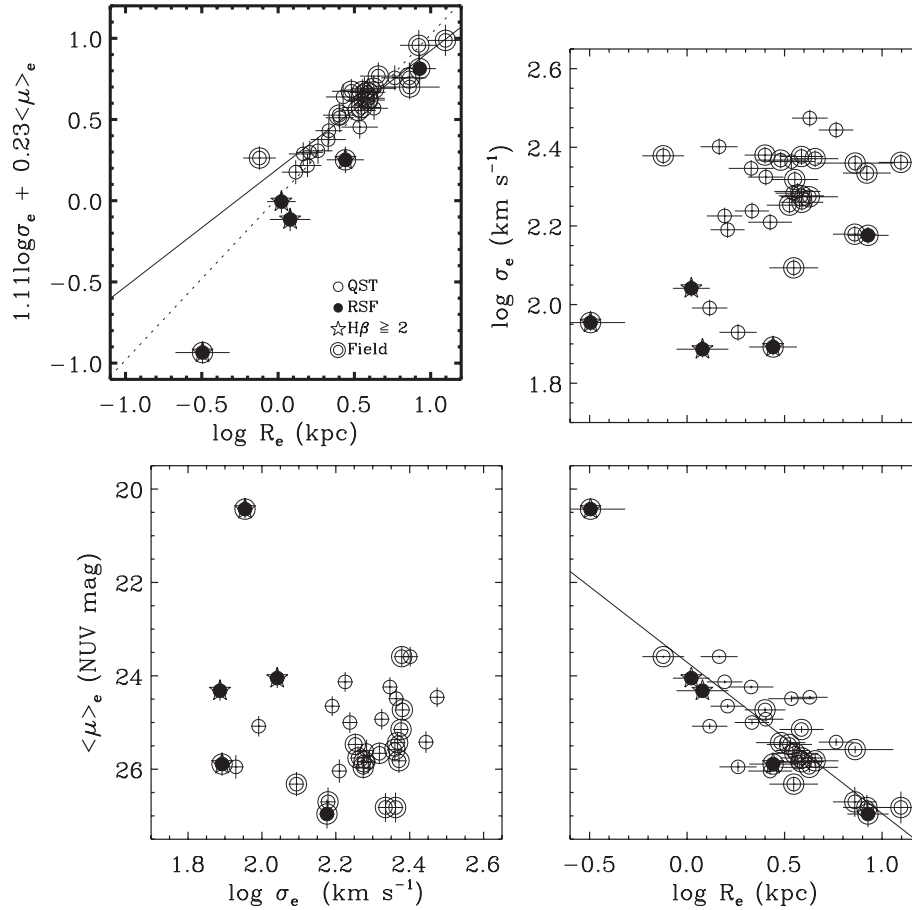


Figure 6. FP and projections of the FP in the FUV band. Symbols are the same as in Fig. 3 and linear fits to the two subsamples (whole sample, quiescent galaxies only) are shown as dotted and solid lines, respectively. We also present a linear fit to the whole sample as a solid line in the bottom-right plot (Kormendy relation).

**Figure 7.** Same as Fig. 6 but for the NUV band.**Table 3.** FP coefficients at FUV and NUV.

Band	a	b	Scatter	Population
FUV	1.09	0.22	0.17	All
FUV	1.25	0.29	0.14	QST
NUV	1.11	0.23	0.14	All
NUV	1.82	0.30	0.08	QST

from a least-square fit by minimizing the variations to the planes. Fits including and excluding RSF galaxies are listed. As mentioned above, RSF galaxies tend to have lower velocity dispersions and smaller effective radii and thus smaller masses than the bulk of the sample galaxies. Crucially, RSF galaxies systematically deviate from the best-fitting planes so as to create shallower slopes, and they significantly increase the scatters. This is clearly shown by both Figs 6 and 7 and the FP parameters listed in Table 3. Although our sample is limited, this indicates that lower mass galaxies, forming a larger amount of young stars, more greatly (and systematically) deviate from the plane defined by the quiescent galaxies.

The long-wavelength K -band slope is less sensitive to contamination from young stellar populations than optical bands ($a \approx 1.53$; see Pahre et al. 1998a). Considering our NUV FP fit for the entire sample, our slope is significantly different from both the virial expectation and the K -band value. However, comparing our NUV-derived slope of 1.82 for quiescent galaxies only, we recover all of the K -band slope and almost all of the tilt away from the virial

expectation. We thus conclude that a significant fraction of the FP tilt and scatter is due to low-mass early-type galaxies with stellar populations significantly younger than those of high-mass galaxies, thus reducing their mass-to-light ratios. The FUV FP is more complicated to interpret because the light contribution from the old but hot stars (UV-upturn phenomenon) is significant in the FUV., but a small, if weaker, trend is observed.

The tilt of the FP was also considered in Paper IV using V -band photometry and sophisticated dynamical modelling. The main result to emerge was a clear, continuous and monotonic increase of the total dynamical mass-to-light ratio with velocity dispersion, consistent with the FP tilt. As can be seen here in Fig. 4, the NUV blue galaxies affecting the tilt of the UV FPs also have some of the smallest velocity dispersions and lowest virial masses. Furthermore, although there is non-negligible scatter, clear $(\text{NUV}-V)-\sigma_e$ and $(\text{FUV}-\text{NUV})-\sigma_e$ [or $(\text{NUV}-V)-M_{\text{vir}}$ and $(\text{FUV}-\text{NUV})-M_{\text{vir}}$] relations are observed. The current results thus seem entirely consistent with that of Paper IV. We are simply picking up the galaxies with the largest and/or most RSF events here with our integrated $\text{NUV}-V < 5.0$ criterion.

Many studies have also exploited the projection of the FP on to the plane defined by R_e and $\langle\mu\rangle_e$, i.e. the Kormendy relation (Kormendy 1977). A clear advantage of this relation is that it is based on photometric parameters only, which are easily accessible even at high redshifts. All of our sample galaxies follow the Kormendy relation, and we note that our UV Kormendy relation slopes (3.48 ± 0.49 and 3.24 ± 0.36 at FUV and NUV, respectively; see Figs 6 and

7) are in rough agreement with the optical one ($\langle \mu \rangle_e \approx 3.02 \log R_e$ in Hamabe & Kormendy 1987). It may, however, be significant that, compared to the mean trends, RSF galaxies tend to be slightly offset towards smaller and/or fainter objects.

3.3 UV radial colour profiles

We use the radial colour profiles to discuss the stellar populations of the sample galaxies. We discuss our criteria to classify them in groups exhibiting similar properties below, and then discuss in more detail galaxies with signs of RSF.

First, our main tool are the NUV–V radial colour profiles. As for integrated magnitudes, we classify a galaxy as having had RSF if its NUV–V profile is below 5.0 for all or a significant portion of the radii probed. Galaxies included in this category are NGC 474, 1023, 2974, 3032, 4150, 4459, 4526, 4550 and 7457. Having said that, NUV–optical colours normally show negative slopes (the central regions being redder) due to variations in the properties of the underlying stellar populations, mainly age and metallicity. It is well-known that there is a significant age–metallicity degeneracy (see e.g. Worthey 1994), whereby it is difficult to distinguish the effects of a small change in age from those of a small change in metallicity. Therefore and even though we recognize the qualitative aspect of such a criterion, if only a small part – usually outside the effective radius – of the UV–optical colour profiles goes smoothly and only a little below the NUV–V = 5.0 cut, we consider it as having been caused by a simple metallicity gradient rather than residual star formation. In other words, localized regions with NUV–V < 5.0 are only classified as recently star-forming if the changes in colour are abrupt. The galaxies exhibiting this behaviour include NGC 821, 2699, 4546, 4570, 5308, 5813, 5831 and 5845. Their integrated colours are all redder than the NUV–V = 5.0 colour threshold.

Secondly, we consider the FUV–NUV profiles. If a galaxy without evidence of RSF has a central region with FUV–NUV < 1.0, we identify it as a UV-upturn galaxy. These galaxies include NGC 4374, 4486, 4552, 4621 and 5846. As expected, these galaxies are among the most luminous and most massive in our sample. All except NGC 4621 are slow rotators (see Paper IX). Significantly, some UV-upturn galaxies have extreme FUV–NUV and FUV–V colours at small radii, bluer than the integrated aperture colours of classic UV-upturn galaxies dating from the *International Ultraviolet Explorer* era. As these galaxies provided important constraints on stellar evolutionary models of the UV upturn, it will be interesting to revisit the models with the current data. It is also probably worthwhile to point out that a number of galaxies exhibit evidence for both a UV upturn (FUV–NUV < 1.0 in the centre) and a substantial age or metallicity gradient (NUV–V < 5.0 in the outer parts). These galaxies are NGC 2695, 4278, 4564, 5198 and 5982. They probably harbour milder forms of the UV-upturn phenomenon, where neither the UV upturn nor large-scale stellar population gradients dominate.

Finally, we note that some galaxies exhibit no evidence for RSF or a UV upturn, and no substantial large-scale age and/or metallicity gradient. These galaxies are NGC 524, 2768, 4387, 4458, 4473, 4477 and 5838.

Of the nine galaxies with convincing evidence for RSF, five (NGC 474, 1023, 2974, 4459 and 4526) show blue UV–optical colours in some regions only (as opposed to all radii). Interestingly, this appears connected to kinematic (sub)structures detected in SAURON and other studies. The case of NGC 474, known

for its shells (e.g. Turnbull, Bridges & Carter 1999) and strong misalignment between the kinematic and photometric major axes (see Paper III), shows blue UV–optical colours and enhanced H β line strength at large radii. NGC 1023, with a prominent twist in the centre of the velocity map (Paper III), reveals a UV blob on the eastern side of the galaxy and blue UV–optical colours in the same region, probably related to an interacting gas-rich dwarf galaxy (see Morganti et al. 2006 for H I observations). The presence of a UV-prominent outer ring related to a bar in NGC 2974 was discussed by Jeong et al. (2007). Blue UV–optical colours in this galaxy are observed in the central and outer regions. In the case of NGC 4459, blue UV–optical colours are observed only in the centre. Paper VI shows that this galaxy has strong H β absorption within a central dust ring associated with a decoupled stellar and ionized gas disc (Papers III and V, respectively). NGC 4526 is a similar object appearing more edge-on. Both NGC 4459 and 4526 also harbour central molecular gas discs (Young et al. 2008).

On the other hand, of the nine galaxies with RSF, NGC 3032, 4150, 4550 and 7457 show *overall* blue UV–optical colours. NGC 3032 shows the bluest UV–optical colour in our sample and has a prominent peak in the H β map suggesting recent and ongoing star formation. McDermid et al. (2006) identified two types of kinematically decoupled components (KDCs) among the SAURON sample: compact KDCs, which are often young and occur in fast-rotating galaxies, and kiloparsec-scale KDCs, with homogeneously old populations in slow-rotating galaxies (see also Paper IX). NGC 3032, 4150 and 7457 all have small-scale KDCs (Paper III) and are considered fast rotators (Paper IX). NGC 3032 and 4150 further have substantial molecular gas discs (Young et al. 2008). In the case of NGC 4550, with two counter-rotating stellar discs (see e.g. Rubin, Graham & Kenney 1992; Rix et al. 1992), the H β line strength is elevated in the central region along the major axis (Paper VI). Its molecular gas content and detailed SFH are discussed in Crocker et al. (2008). NGC 7457 has a controversial CO detection discussed previously (see Welch & Sage 2003; Combes et al. 2007).

4 SUMMARY

We have presented *GALEX* FUV and NUV imaging along with ground-based F555W imaging from the MDM Observatory for 34 early-type galaxies from the SAURON survey sample. Nine of them show extended blue UV–optical colours, hinting at RSF. Five of these are also classified as RSF early-type galaxies by our integrated CMR criterion. Supporting the findings from the UV CMR technique, four out of the five candidate RSF galaxies also show enhanced H β absorption line strengths, a widely used post-starburst signature. The only exception (NGC 474) is noted for its outer shells. Roughly 15 per cent of the early-type galaxies in our sample are therefore classified as RSF early type galaxies. Considering that the UV flux from a starburst is only detectable for roughly 1 Gyr, this implies that residual star formation has been common in early-type galaxies even up to the present day. This star formation interpretation is also consistent with the results of molecular gas surveys, as the CO detection rate is roughly 60 per cent in RSF early-type galaxies, much higher than in quiescent systems.

The velocity dispersions of these galaxies with evidence of RSF are among the lowest in the sample, and none has $\sigma_e > 200 \text{ km s}^{-1}$, consistent with the independent work of Schawinski et al. (2007). RSF early type galaxies also tend to have smaller effective radii and thus smaller (virial) masses. Despite the limited number statistics, a key result from the present study is that RSF early type galaxies

change the slopes of scaling relations (CMRs and FPs) and dominate the scatters in them. Most notably, much of the FP tilt and scatter can now be explained by the fact that the properties of a substantial fraction of early-type galaxies are influenced by RSF, systematically biasing their mass-to-light ratios, especially at low masses. The UV FPs become significantly closer to the virial expectation (and tighter) when only quiescent early-type galaxies are considered. The same must be true at optical wavelengths, although the effect will be smaller. Our result appears consistent with the $M/L-\sigma_e$ relation derived in Paper IV, although the current colour threshold only picks up the most extreme objects.

The radial UV–optical colour profiles not only reveal RSF galaxies, but also a number of galaxies with smooth large-scale age and/or metallicity gradients. Similarly, a number of galaxies exhibiting the UV-upturn phenomenon are identified. Some have bluer FUV–NUV and FUV–V colours at small radii than the integrated aperture colours of classic UV-upturn galaxies. A number of galaxies show both a central UV upturn and large-scale gradients. This diversity of behaviours will, in due time, need to be explained and reproduced by stellar population models.

Early-type galaxies are no longer thought to be simple. Many of them have kinematic anomalies and substructures that are non-trivial to interpret. We have shown here that they are also composed of multiple generations of stars widely separated in age. This is consistently found by other short-wavelength as well as far-infrared studies, mainly aided by space experiments. The presence of young stars in seemingly old populations is no longer debatable, yet it is still unclear what causes and regulates this residual star formation. We believe that our data base, spatially resolving nearby early-type galaxies, will be a key to answering some of those questions.

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APPENDIX A: DESCRIPTION FOR INDIVIDUAL GALAXIES

We briefly comment here on the *GALEX* and MDM surface brightness profiles of the E/S0 galaxies presented in this paper. Individual comments on the stellar kinematics, the emission-line maps and the line-strength structures are given in Papers III, V and VI, respectively.

NGC 474. This galaxy (Arp 227), noted for its shell structures (e.g. Turnbull et al. 1999), shows blue NUV–V colours except in the central regions. The H β absorption line strength rises towards larger radii (Paper VI). Considering the UV–V colours, it seems that young stellar populations are present in the outer parts.

NGC 524. This galaxy shows red colours in both the FUV–V and the NUV–V colour profile, making a good example of a quiescent elliptical galaxy.

NGC 821. This edge-on galaxy with a rapidly rotating disc-like component (Paper III) shows red FUV–NUV and UV–V colours except outside the effective radius.

NGC 1023. This SB0 galaxy shows a tilt in the FUV and NUV surface brightness profiles around 25 arcsec, a central concentration in all metal lines (Paper VI) and significantly negative values of both h_3 and h_4 (Paper III). We discovered a strong NUV blob, constituting strong evidence for recent/ongoing star formation, just outside the galaxy. This is likely the result of a tidal interaction between NGC 1023 and a neighbouring gas-rich galaxy (see Morganti et al. 2006).

NGC 2695. This galaxy shows blue FUV–NUV colours in the central regions with a high amplitude in h_3 . The UV–V colours become bluer towards larger radii, albeit with large uncertainties. We thus consider this a red galaxy.

NGC 2699. The mean velocity map of this galaxy shows a rapidly rotating component around 5 arcsec (Paper III). Similar to NGC 2695, it shows moderately blue NUV–V colours only in the outer parts.

NGC 2768. This galaxy with a cylindrical velocity field shows red colours in the UV–V colour profiles.

NGC 2974. This galaxy with a UV outer ring shows a blue colour in the outer regions, suggesting young stellar populations associated with a large-scale bar (see Jeong et al. 2007), although the H β line-strength map appears relatively flat (Paper VI).

NGC 3032. Blue UV–V colours are observed everywhere, suggesting a strong ongoing starburst. This galaxy also shows strong H β absorption with negative Mgb and Fe5270_o line-strength gradients (Paper VI).

NGC 4150. This galaxy with a counter-rotating core (Paper III) shows blue UV–V colours and enhanced H β absorption with a strong drop in the Mgb line strength (Paper VI).

NGC 4278. This galaxy shows blue FUV–NUV colours within 20 arcsec, suggesting a strong FUV flux. It also has the strongest line emission (Paper V).

NGC 4374 (M84). This giant elliptical galaxy shows a similar FUV–NUV colour trend as that in NGC 4278. It is known for its BL Lac nucleus and shows strong ionized gas emission (Paper V).

NGC 4387. This boxy galaxy shows red UV–V colours.

NGC 4458. The FUV image of this galaxy exhibits an indistinct shape except for the central regions. This faint disc-like object also has a small central kinematically decoupled core (Paper III).

NGC 4459. This galaxy with a central dust ring (Paper V) shows blue UV–V colours within 10 arcsec, coincident with an intense H β absorption line strength and molecular gas.

NGC 4473. This galaxy with a high-velocity dispersion along the major axis likely resulting from substantial counter rotation (Paper III) shows extended UV emission in the outskirts.

NGC 4477. This galaxy with a prominent misalignment between the stellar kinematic and photometric major axes (Paper III) shows overall red UV–V colours.

NGC 4486 (M87). This galaxy with a prominent jet shows blue UV–V colours in the central regions, almost certainly related to nuclear activity. The jet also induces large errors in the surface brightness profiles.

NGC 4526. The UV–V plots of this galaxy exhibit blue colours in the centre, coincident with strong H β absorption (Paper VI) and a stellar velocity dispersion drop due to a star-forming fast-rotating disc (Papers III and V).

NGC 4546. This galaxy shows moderately blue NUV–V colours in the outer parts. These could be related to the likely presence of a bar and associated star formation.

NGC 4550. This galaxy with two counter-rotating stellar discs (see Rubin et al. 1992; Rix et al. 1992) shows blue UV–V colours everywhere.

NGC 4552 (M89). This famous UV-upturn galaxy shows very blue FUV–NUV colours in the centre.

NGC 4564. This elongated galaxy with a disc-like component (Paper III) shows blue FUV–NUV colours within the inner 10 arcsec.

NGC 4570. This fast-rotating edge-on galaxy shows slightly blue UV–V colours in the outer parts. The presence of a bar was claimed by van den Bosch & Emsellem (1998).

NGC 4621. This object has a disc-like component similar to NGC 4564 (Paper III). A counter-rotating component is also detected inside 2 arcsec (Wernli, Emsellem & Copin 2002). Blue FUV–NUV colours are observed in the centre.

NGC 5198. This galaxy with a central kinematically decoupled core shows blue FUV–NUV colours in the centre that may indicate the presence of a UV upturn.

NGC 5308. This edge-on disc galaxy shows red UV–V colours.

NGC 5813. This galaxy with a well-known kinematically decoupled core shows marginally blue NUV–V colours in the outer parts.

NGC 5831. Like NGC 5813, this galaxy with a kinematically decoupled core shows blue NUV–V colours in the outer regions. There is a small bump in both FUV and NUV surface brightness profiles around 35 arcsec.

NGC 5838. This fast-rotating object shows red NUV–V colours throughout.

NGC 5845. This compact elliptical galaxy shows a small bump in the NUV profile at a radius of ≈ 10 arcsec, with a corresponding distinct peak in the velocity map related to the central disc.

NGC 5846. Similar to NGC 4552 and 4564, this bright giant elliptical galaxy shows blue FUV–NUV colours in the centre that are suspected to be a UV upturn, with marginally blue NUV–*V* colours.

NGC 5982. This galaxy with a kinematically decoupled core shows blue FUV–NUV colours in the centre.

NGC 7457. This kinematically decoupled core galaxy shows blue NUV–*V* colours throughout, suggesting RSF. The $H\beta$ absorption is also high everywhere (Paper VI).

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